

Possible Signature of Sausage Waves in Photospheric Bright Points

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Abstract Sausage waves have been frequently reported in solar magnetic structures such as sunspots, pores, and coronal loops. However, they have not been unambiguously identified in photospheric bright points (BPs). Using high-resolution TiO image sequences obtained with the *Goode Solar Telescope* at the Big Bear Solar Observatory, we analyzed four isolated BPs. It was found that their area and average intensity oscillate for several cycles in an in-phase fashion. The oscillation periods range from 100 to 200 seconds. We interpreted the phase relation as a signature of sausage waves, particularly slow waves, after discussing the sausage wave theory and the opacity effect.

Keywords: Photospheric Bright Points; Waves, Magnetohydrodynamic; Magnetic Flux Tube; Opacity Effect

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1. Introduction

Photospheric bright points (BPs) are dynamic, small-scale, bright structures in the solar photosphere, usually appearing in dark intergranular lanes. They are associated with magnetic elements in magnetograms and have a **magnetic-field** strength of about **1 kG**. BPs are typically circular (Berger et al., 1995; Bovelet and Wiehr, 2003), while some of them have elongated shapes (Kuckein, 2019). Their equivalent diameter is about **100–300 km**, and their lifetime is about **90 seconds–10 minutes** (Berger and Title, 1996; Utz et al., 2010; Keys et al., 2011, 2014). In addition, BPs usually have a brightness contrast of **0.8–1.8** relative to the mean intensity of the photosphere (Sánchez Almeida et al., 2004).

BPs were first discovered in G-band observations of the photosphere in the 1970s (Dunn and Zirker, 1973), and later Mehlretter (1974) found that they represent **magnetic-flux** concentrations by comparing Ca II K images with magnetograms. BPs are therefore also called magnetic bright points, and they are considered as the photospheric footpoints of magnetic flux tubes (e.g. Berger et al., 1998; De Pontieu, 2002). A process named convective collapse is believed to be their formation mechanism (Spruit, 1979), which has received some observational support (e.g. Solanki et al., 1996; Bellot Rubio et al., 2001; Nagata et al., 2008).

Since BPs can be regarded as tracers of photospheric magnetic flux tubes, many studies have suggested that BPs may host a variety of MHD waves (Martínez González et al., 2011; Jess et al., 2012a; Mumford, Fedun, and Erdélyi, 2015; Jafarzadeh et al., 2017). Given that the dissipation of MHD waves is among the most promising mechanisms for heating the solar corona, it is of great significance to study MHD waves associated with BPs or photospheric magnetic elements. So far, possible observational manifestations of kink waves (Stangalini, Berrilli, and Consolini, 2013; Stangalini, Giannattasio, and Jafarzadeh, 2015) and Alfvén waves (Jess et al., 2009) have already been found in BPs (or magnetic elements). However, sausage waves have not been unambiguously identified in them.

Sausage waves (or sausage modes) in a flux tube are characterised by a sausage-like boundary of the tube, first defined by Defouw (1976). A detailed discussion of this wave mode is given in **Section 4.1**. There are many studies focusing on theoretical understanding and seismological applications of sausage modes (e.g. Chen et al., 2015; Guo et al., 2016; see also the review by Li et al., 2020). In observations, sausage modes in **the** photosphere have been widely studied in magnetic pores and sunspots (e.g. Dorotovič, Erdélyi, and Karlovský, 2008; Morton et al., 2011; Grant et al., 2015; Freij et al., 2016). Also, Morton et al. (2012) found evidence for the simultaneous existence of sausage and kink modes in fine magnetic structures in the chromosphere. In the corona, short-period, quasi-periodic pulsations (QPPs) in flares are also generally interpreted as sausage waves (e.g. Nakariakov, Melnikov, and Reznikova, 2003; Van Doorselaere et al., 2011; Tian et al., 2016; see also the review by Zimovets et al., 2021).

In this **article**, we report the existence of in-phase oscillations of area and average intensity in BPs using high-resolution TiO images for the first time. Our findings can be regarded as possible evidence of sausage modes in **photospheric**

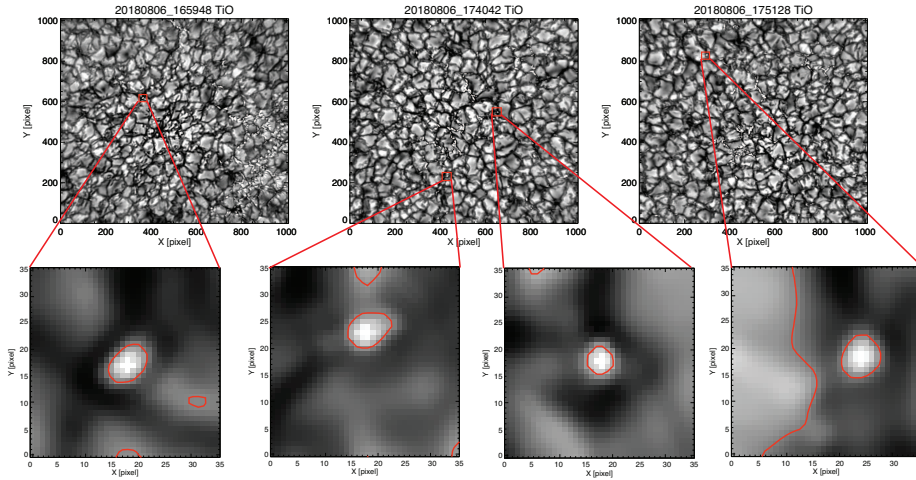


Figure 1. TiO images with BPs marked by boxes. The images below show **an enlarged** view of the red box regions, where BPs are identified and tracked with red contours.

flux tubes. In fact, magnetoacoustic waves characterised by intensity or line-of-sight (LOS) velocity oscillations have been found in BPs before (Jess et al., 2012a,b; Stangalini et al., 2013; Jafarzadeh et al., 2017). But these previous studies could not tell whether these waves **were** sausage waves, because either they did not consider BP areas or the area variations were too weak. Martínez González et al. (2011) studied four magnetic patches in circular polarization images obtained by **the Imaging Magnetograph eXperiment (IMaX) on board *Sunrise***, and **they** found that the iso-flux area oscillates with time. However, the time-varying periods suggested that the oscillations may be caused by granular motions rather than a characteristic oscillation mode. Fujimura and Tsuneta (2009) also focused on oscillations of intergranular magnetic structures (IMs) in a magnetogram and interpreted these oscillations as sausage and/or kink waves. But the lifetimes of these IMs are all **greater than** one hour, **which is** much longer than the typical lifetime of BPs. Kolotkov et al. (2017) also detected long-period (**80 – 230 minutes**) quasi-periodic oscillations of the average magnetic field and total area of a small magnetic **structure**. More recently, Keys et al. (2020) studied the rapid variations of magnetic fields in BPs and believed that such changes are consistent with the behavior of MHD waves, but they did not discuss in detail the changes of intensity.

This **article** is organised as follows. In Section 2, we describe our observation and data analysis methods, and the key results are presented in Section 3. We discuss the **wave-mode** identification in Section 4 before giving the final conclusion and some remarks in Section 5.

Table 1. Observational details of the BPs. The area is averaged over the time, and C.C. represents the correlation coefficient of average intensity and area.

	Date	Time [UT]	Location	Area [Mm ²]	C.C.
BP1	06 Aug. 2018	16:54:18–17:03:18	(501'', –222'')	0.015	0.71
BP2	06 Aug. 2018	17:38:12–17:46:28	(–260'', 2'')	0.017	0.92
BP3	06 Aug. 2018	17:33:12–17:41:27	(–260'', 2'')	0.013	0.69
BP4	06 Aug. 2018	17:47:43–17:53:43	(–260'', 2'')	0.015	0.89

2. Observations and Data Analysis

The data used in this study were acquired with the 1.6-meter *Goode Solar Telescope* (GST; Cao et al., 2010) at the Big Bear Solar Observatory (BBSO). We used the Broadband Filter Imager (BFI) installed on the GST to observe the TiO band (7057 Å) from 16:28 to 17:59 UT on 6 **August** 2018, with a cadence of 15 **seconds**. The observed regions include the active region NOAA 12717 at (501'', –222'') from 16:28 to 17:33 and the western lobe of a coronal hole at (–260'', 2'') from 17:33 to 17:59.

The bandwidth of the filter is 10 Å. The **field of view** (FOV) of the TiO data was 77'' × 77'' and the pixel size was 0.034''. During the observation, the seeing condition was better than 3'', which is good for the identification of BPs. Speckle reconstruction and a high-order adaptive optics (AO) system were used to reach the **spatial-resolution** limit. The data were also destretched to remove the residual **atmospheric-seeing** effects.

We first selected isolated BPs in the FOV by **eye**. Liu et al. (2018) and Kuckein (2019) divided BPs into isolated ones, which are individual, and non-isolated ones, which display a splitting or merging behavior in observations. Obviously, it is more suitable to use isolated BPs to search for MHD waves. **Because for non-isolated BPs, their splitting and merging behaviors may affect our identification of waves.** For the selected BPs, we used contours with a fixed intensity threshold level to identify and track them. Then we calculated the area and average intensity of these BPs at each time and examined if there is any oscillation signal. Finally, four of the BPs are found to be well isolated and there is also possible evidence of sausage waves inside. Their positions and shapes are shown in Figure 1. We named them BP1 to BP4. When calculating their average intensity and area, we carefully removed the unrelated contours nearby (we can see them around BP1, BP2, and BP4 in Figure 1) by manually zooming **on** the region of interest for every frame. We summarize the properties of the four BPs in Table 1.

3. Result

Figure 2 presents the time series of the area and average intensity of these four BPs, and we can see obvious oscillations for **two to four** cycles. It can also

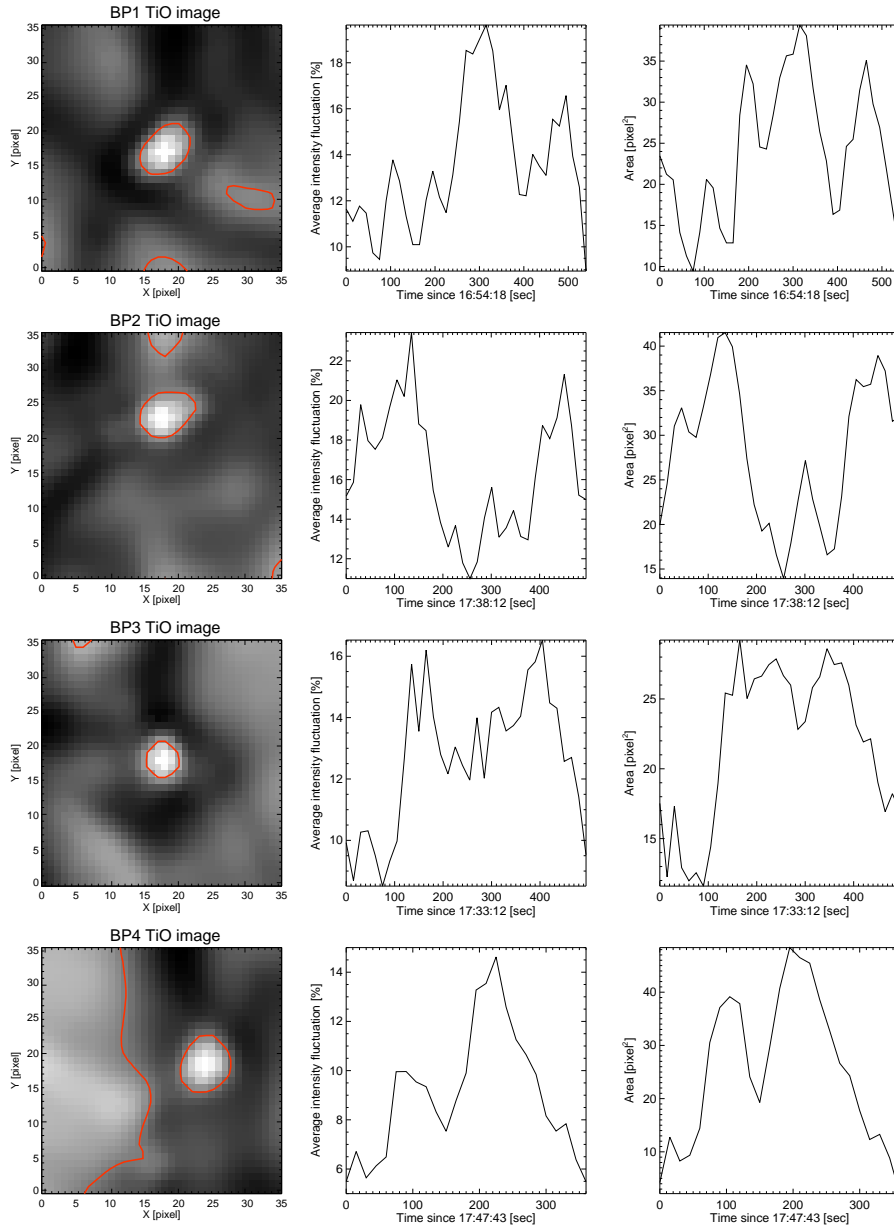


Figure 2. Time series of average intensity and area of the four BPs. Each row corresponds to one BP, and the right two columns show the temporal variations of the average intensity and area, respectively. Here we use the intensity fluctuation percentage with respect to the mean value of the background. (**Animation** of this figure is available in the **Supplementary Materials**.)

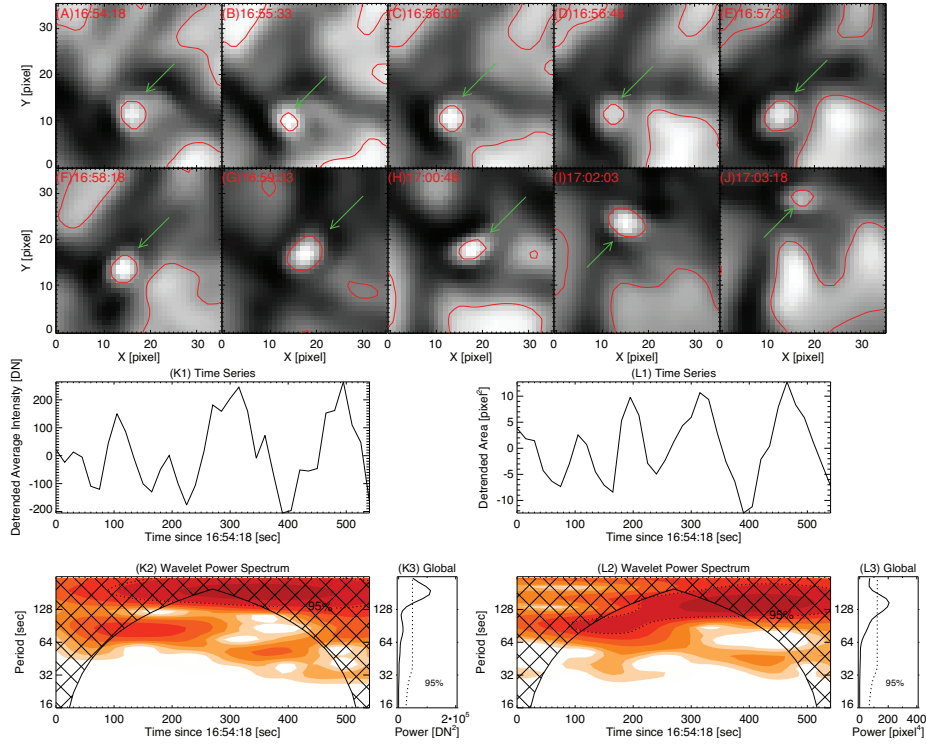


Figure 3. (A – J): snapshots of BP1 marked by green arrows at **ten** different time. Approximately, every three adjacent panels correspond to an oscillation cycle. (K): wavelet analysis results of detrended time series (shown in Panel K1) of average intensity. The results include the wavelet power spectrum (Panel K2) and the global power spectrum (Panel K3). Darker colors represent stronger powers. The dashed lines indicate a significance level of 95%. (L): same as Panel K but for the detrended time series of area.

be noticed that for each BP, the crests and troughs for these two **features** appear at nearly the same time, which suggests that the oscillations of area and average intensity are in phase. Among these BPs, BP1 has the most oscillation cycles, with four clear crests in the time series of area, respectively at **105, 195, 315, and 465 seconds**. These instants correspond well to crests in the curve of average intensity. **In addition**, there is also a good correspondence between moments of each trough. Figure 3 gives ten snapshots of BP1 in **Panels A – J**, each corresponding to a moment of wave crest or trough. We can clearly see that the BP’s area oscillates with time.

Then we detrended the time series of average intensity and area by subtracting a **three-minute** running average before performing wavelet analysis. The detrended series and wavelet analysis results are shown in Panels K and L of Figure 3. The peak global wavelet powers for the average intensity and area are at around **150 seconds**, which could be seen as the oscillation period. **For** the other BPs, the periods can be estimated to be **100 – 200 seconds**. The period range we found here are comparable to the results in some previous studies, which interpreted the oscillations they found in BPs as magnetoacoustic

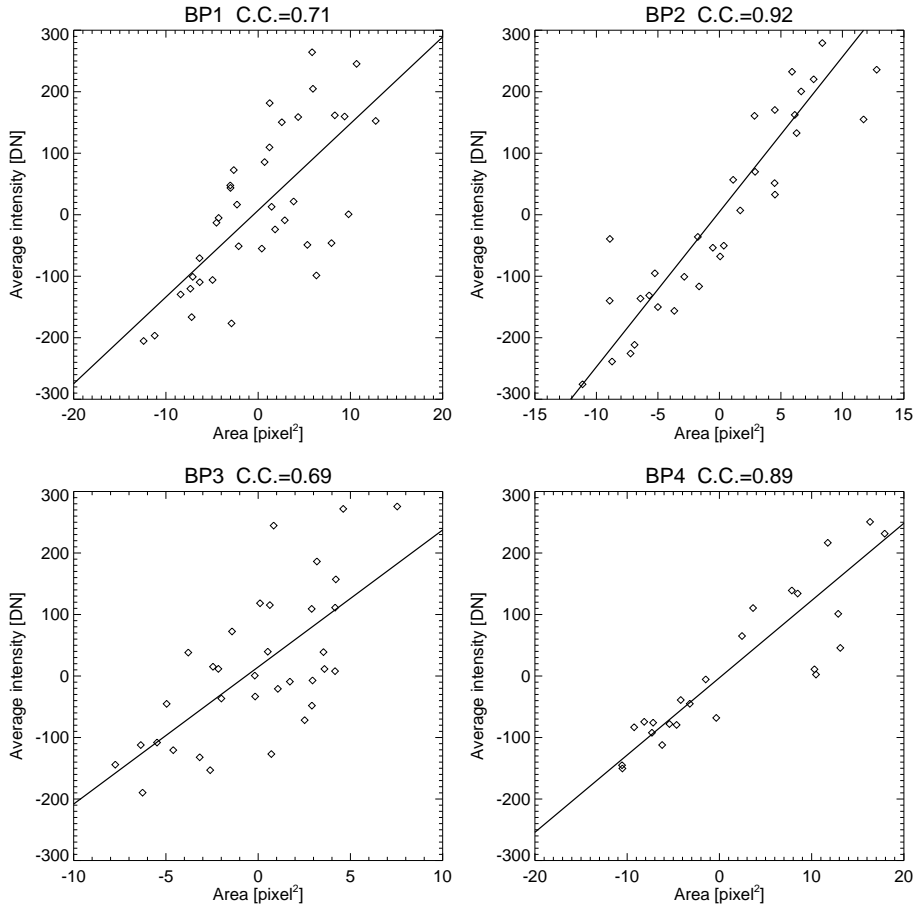


Figure 4. Scatter plots of the relationship between average intensity and area for the four BPs, with solid lines representing linear fits. Note that the average intensity and area are both detrended.

waves (Jess et al., 2012a,b). We note that oscillations with longer periods above **four minutes** were detected in small magnetic structures before (Fujimura and Tsuneta, 2009; Martínez González et al., 2011; Kolotkov et al., 2017).

In order to further confirm the in-phase relation between oscillations of average intensity and area, we drew scatter plots of relationship between them, and calculated the correlation coefficients, as shown in Figure 4. It can be seen that the correlation coefficients are all above 0.68. For BP2 and BP4, the coefficients are around 0.9, showing a good **positive** correlation. Meanwhile, we also performed a lagged cross-correlation (LCC) analysis. From Figure 5, we can see that the coefficients for all BPs reach the peak when time lag is **zero**. All the results confirm the in-phase relation between average intensity and area.

A possible question is whether the correlation or the in-phase relation is caused by our analysis method, i.e. using contours with a fixed intensity threshold to identify BPs and calculate their parameters. If we consider the total intensity

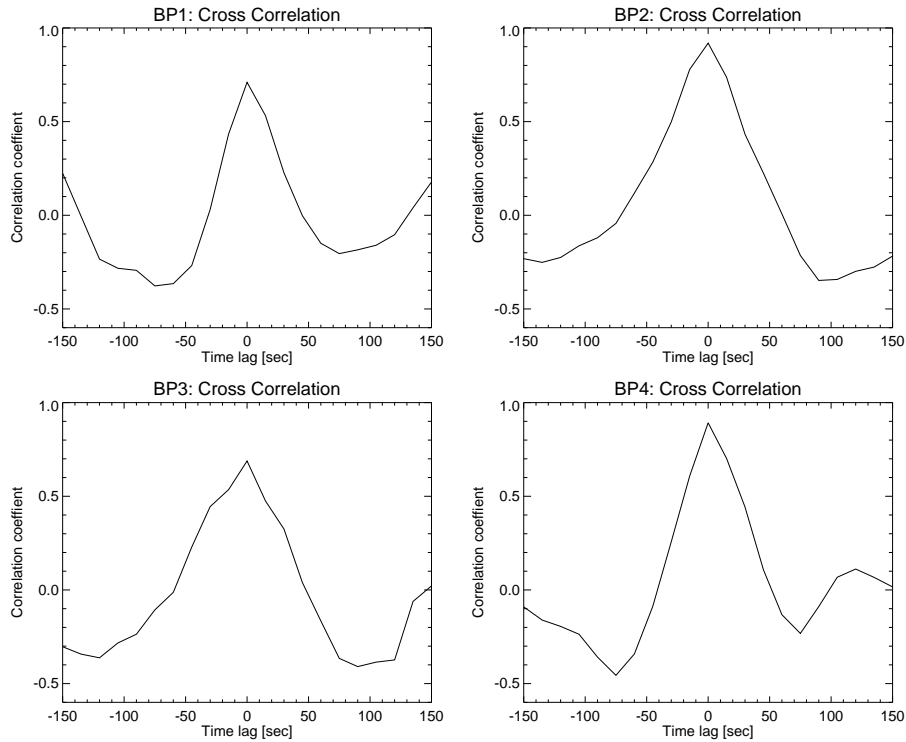


Figure 5. Lagged cross-correlations between average intensity and area. The time lag is between the detrended time series of area and average intensity.

rather than the average intensity in the contours, we might tend to obtain an in-phase relation, since it seems that the total intensity will always increase (or decrease) along with the increase (or decrease) of the area. So here we chose to use time series of the average intensity. **Additionally**, we studied for more than **four** isolated BPs with the same method, and **we** found that there are cases without such an oscillation pattern. In conclusion, the in-phase relation is not a result of our **data-processing** method, but more likely associated with some characteristic oscillation modes.

We will show in the next section that the observed in-phase oscillations might be explained as sausage waves in BPs.

4. Discussion

4.1. Sausage Mode Waves in a Flux Tube

In the solar atmosphere, magnetic flux tubes are formed with magnetic field lines clumping into tight bundles. They can be regarded as channels guiding MHD waves. A flux tube is usually assumed as a straight cylinder described by cylindrical coordinates (r, ϕ, z) . The tube boundary can oscillate in different

geometrical patterns, forming various wave modes including sausage modes, kink modes, fluting modes, and torsional (Alfvén) modes (Roberts, 2019). For sausage modes, the axis of the tube is motionless, while the cross-section area fluctuates with time, resulting in a sausage-like axisymmetric perturbation of the tube boundary. So sausage waves are axisymmetric waves, belonging to magnetoacoustic waves. According to the relationship between phase speed c_p and sound speed c_s in the tube, we can divide the sausage **wave** into **a** fast mode and slow mode. The former has a faster phase speed than **the** sound speed, while the latter has a slower phase speed. Assumed here and hereafter is that the characteristic speeds inside and outside the tube satisfy the ordering for “photospheric conditions” in the sense of **Figure 3** of Edwin and Roberts (1983). In particular, β inside the tube is assumed to be smaller than unity. This assumption on the internal plasma β is generally accepted for photospheric BPs, **even though** β is difficult to **measure directly** (e.g. Shelyag et al., 2010; Cho et al., 2019).

Based on ideal MHD equations, **setting the radial-velocity** perturbation $v_r = A(r) \cos(kz - \omega t)$, we can obtain (Freij et al., 2016, without considering azimuthal components of **the** velocity perturbation and magnetic field):

$$b_z = \frac{B_0}{\omega} \frac{1}{r} \frac{\partial(rA(r))}{\partial r} \sin(kz - \omega t), \quad (1)$$

$$p_1 = c_s^2 \rho_1 = -\frac{\omega \rho_0 c_s^2}{(c_s^2 k^2 - \omega^2)} \frac{1}{r} \frac{\partial(rA(r))}{\partial r} \sin(kz - \omega t), \quad (2)$$

where b_z is the longitudinal magnetic-field perturbation, B_0 and ρ_0 are background magnetic field and density, p_1 is the pressure perturbation, and ρ_1 is the density perturbation.

Comparing Equations 1 and 2, we can **state** that if $c_s^2 k^2 - \omega^2 > 0$ (i.e. the phase speed $c_p = \omega/k$ is smaller than the sound speed c_s , corresponding to slow sausage mode), then b_z and ρ_1 (or p_1) will be 180° out of phase (i.e. in anti-phase). On the contrary, if $c_s^2 k^2 - \omega^2 < 0$ (i.e. the phase speed is larger than the sound speed, corresponding to **the** fast sausage mode), then the perturbations will be in phase.

Since the magnetic flux in a flux tube is constant, we have $B_0 S_0 = (B_0 + b_z)(S_0 - S_1)$. Here, S_0 and S_1 represent the unperturbed and perturbed cross-section areas, respectively. Ignoring $b_z S_1$, we can obtain $B_0 S_1 = -b_z S_0$, or

$$\frac{S_1}{S_0} = -\frac{b_z}{B_0}. \quad (3)$$

So the area perturbation S_1 and the longitudinal field perturbation b_z are in anti-phase, and we can **find** the change of **the** magnetic field from the change of the area. This was also used **by** Grant et al. (2015).

As a result, for **the** slow sausage mode, fluctuations of area and density are in phase; whereas for **the** fast sausage mode, the fluctuations are in anti-phase. These simple phase relations have also been confirmed by other theoretical

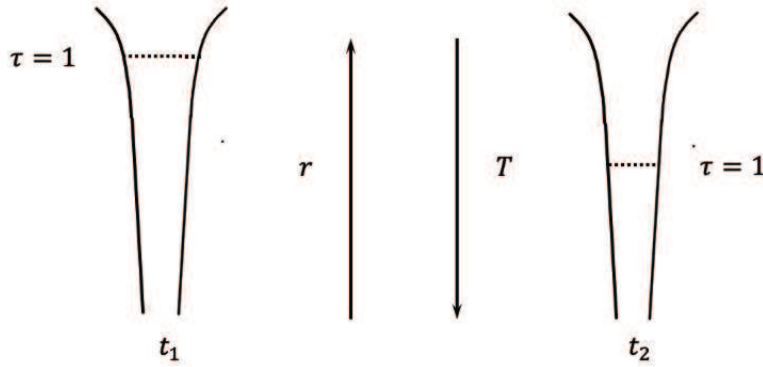


Figure 6. The change of intensity and area caused by the opacity effect in a flux tube. The dashed lines represent the $\tau = 1$ surface, corresponding to the BP **that** we see in the TiO image. From time t_1 to t_2 , the surface moves to a lower height. The two arrows in the middle indicate that the height increases from the bottom up and temperature increases from the top down.

investigations (Moreels, Goossens, and Van Doorselaere, 2013; Moreels and Van Doorselaere, 2013).

Next we will discuss the impact of the opacity effect and associate density with the continuum intensity, so that we can obtain the phase relation between fluctuations of intensity and area for fast and slow sausage modes.

4.2. Phase Relation between Fluctuations of Intensity and Area

The TiO data used in our study could be regarded as continuum images (see, e.g. LeBlanc, 2010), so the intensity **that** we observed could be written as

$$I = \int_0^\tau \frac{\sigma T(\tau)^4}{\pi} e^{-\tau} d\tau, \quad (4)$$

where σ is the Stefan–Boltzmann constant, and $T(\tau)$ is the local temperature at optical depth of τ . Thus the observed intensity fluctuations come primarily from two sources: the change of temperature and the change of optical depth. The latter depends on density and temperature **along** the **line-of-sight**, which is known as the opacity effect (Fujimura and Tsuneta, 2009).

If the oscillations of intensity and area in BPs are caused by the opacity effect, then the oscillations will be in anti-phase. In Figure 6, a BP is shown as a laterally expanding flux tube, and the layer of $\tau = 1$ is marked with the dashed line. Suppose that from time t_1 to t_2 , the $\tau = 1$ layer moves to a lower height, resulting in a reduced area of the BP. In the deeper layer with a higher temperature, the intensity would increase, which means that the intensity and area should be in anti-phase. Clearly it is contrary to our observation.

Without considering the opacity effect, we can assume that the **surfaces that** we see are at the same height, as shown in Figure 7. Another assumption **that** we made is that the BP intensity increases from time t_1 to t_2 , and it is caused by the increase of temperature. The higher temperature will lead to a

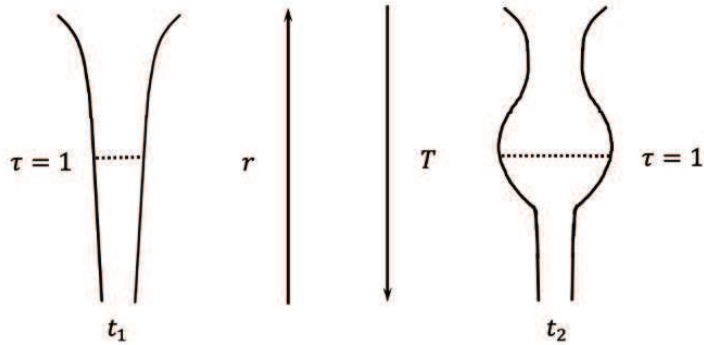


Figure 7. The change of intensity and area caused by the sausage waves. The $\tau = 1$ surface remains the same height from time t_1 to t_2 . The two arrows in the middle indicate that the height increases from the bottom up and temperature increases from the top down.

decrease in the absorption coefficient κ_ν at a frequency ν , which is proportional to $[1 - \exp(-h\nu/k_B T)]$, where h and k_B are the Planck constant and Boltzmann constant (Rutten, 1995). In order to ensure

$$\tau = \int_{r_0}^{\infty} \kappa_\nu \rho(r) dr = 1, \quad (5)$$

the density should increase. Here $\rho(r)$ represents the density at height of r , and r_0 represents height of the surface **that** we see. Our observation shows that the fluctuations of intensity and area are in phase. If the area increases at the same time, the density will be in phase with the area, which is a characteristic of **the** slow sausage wave (see Section 4.1). So the in-phase oscillations of intensity and area that we have observed could be possible evidence of slow sausage waves in BPs.

However, if we consider both the opacity effect and the presence of sausage modes, the result will be different. In Figure 8, from t_1 to t_2 , the sausage waves enlarge the cross-section area of the flux tube, and the magnetic field will decrease. If the opacity effect moves the $\tau = 1$ layer to a lower height at the same time, we also expect in-phase oscillations of intensity and area, but now we cannot distinguish which mode the sausage waves are. The total pressure equilibrium at the same height is given by

$$p_i + \frac{B_i^2}{2\mu_0} = p_e + \frac{B_e^2}{2\mu_0}, \quad (6)$$

where p_i and p_e are the internal and external thermal pressure, B_i and B_e are the internal and external magnetic field, and μ_0 is the permeability. Equation 6 indicates that the decreasing magnetic field inside the tube will lead to an increasing thermal pressure, which is proportional to ρT . If the density ρ increases, the waves will be slow mode; otherwise they will be fast mode. For specific determination, it may be necessary to introduce some appropriate photospheric

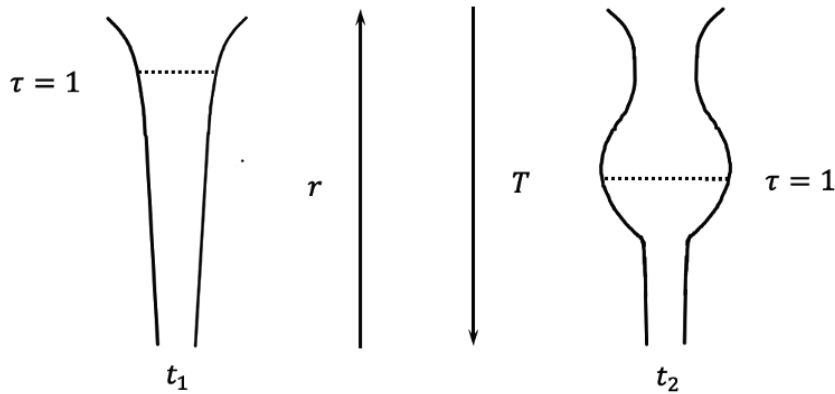


Figure 8. The changes of intensity and area caused by both the opacity effect and sausage waves. The $\tau = 1$ surface moves to a lower height from time t_1 to t_2 . The two arrows in the middle indicate that the height increases from the bottom up and temperature increases from the top down.

atmosphere models (e.g. Fang, Tang, and Xu, 2006), which will be our next step. Nevertheless, no matter which mode the waves actually are, the results **that** we observed can be seen as **evidence** for sausage waves in BPs.

5. Conclusion and Remarks

In this study, we used the TiO images obtained from the GST/BBSO on 6 **August** 2018 to carry out a detailed analysis **of** several isolated photospheric BPs. For the four selected BPs, the changes of their area and average intensity with time show obvious oscillations, and an in-phase relation.

In order to interpret these results, we **presented** a simple theoretical discussion, ruling out the opacity effect alone and identifying the observed oscillations in BPs as sausage modes. Assuming **that** the fluctuations are only caused by sausage waves, we can identify the oscillations in BPs as slow sausage **modes**, which have been found in some previous observations of magnetic pores (e.g. Grant et al., 2015; Freij et al., 2016). Our discovery provides strong support for the existence of sausage waves in the photospheric flux tube. Although the presence of sausage waves has also been observed in magnetic pores or sunspots, BPs are better tracers for the flux tubes, and **they** are more ubiquitous in **the** photosphere. Additionally, BPs are much smaller than pores and sunspots, which makes them more easily buffeted by **surrounding** granular motions. Since the dissipation of sausage modes is generally believed to have an important contribution to the heating of the chromospheric atmosphere (Grant et al., 2015; Yu, Van Doorselaere, and Goossens, 2017; Chen et al., 2018; Keys et al., 2018; Gilchrist-Millar et al., 2021), our findings could be helpful for understanding the chromosphere heating.

Our **future** work will first focus on the mode identification by further addressing the opacity effect. We plan to introduce an appropriate photospheric

atmosphere model to figure it out. **Futhermore**, our research is still limited to single-channel observations of the photosphere. We plan to combine more data in different wavelengths or spectral observations to study the propagation of sausage waves in BPs to the upper atmosphere. Moreover, the *Daniel K. Inouye Solar Telescope* (DKIST: Rast et al., 2021) can provide observations of the photosphere at an unprecedented high resolution. By using the oncoming data from DKIST, we may find more evidence of MHD waves in BPs and carry out more detailed study.

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Disclosure of Potential Conflicts of Interest The authors declare that they have no conflicts of interest.

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